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# A new method of studying the relation between ionization rates and radio-wave absorption in polar-cap absorption events

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**Abstract.** During polar-cap absorption events, which are caused by the incidence of energetic solar protons, the high-latitude ionospheric D region is extended down to relatively low altitudes. While the incoming proton fluxes may be monitored by satellite-borne detectors, and the resulting radio-wave absorption with a ground-based riometer, the enhancement of electron density at a given altitude is less easily determined. Direct measurements by incoherent-scatter radar are infrequent and they tend to lack the necessary sensitivity at the lower levels. Computations of the electron density from the observed particle fluxes are handicapped by uncertainties in the height profile of the effective recombination coefficient.

This paper describes a new approach based on finding the best-fit solution to an over-determined set of equations. The D region is treated as a set of slabs, each contributing to the total radio absorption, and the method relies on the fact that the proton spectrum varies during the event. The analysis produces a set of coefficients relating the absorption increment in the slab to the square root of the production rate, as a function of height. Values of effective recombination coefficient are also deduced over a range of heights, and these agree with previous estimates (Gledhill, 1986) to within a factor of 2. However, whereas the latter do not generally go below 60 km altitude the new determination extends the values down to 40 km.

The new method provides a measurement of the height profile of the absorption in PCA events. It is shown that the slabs centred from 45 to 65 km typically account for 80% of the total daytime absorption, and that less than 1% of the total arises above 80 km or below 30 km. At night most of the absorption comes from the slabs at 75 and 80 km, with no significant contribution from slabs below 75 or above 85 km. These results would not differ significantly from estimates based on the Gledhill profiles if extrapolated downward.

Predictions based on the coefficients generated by the procedure are compared with the polar-cap absorption observed during some recent events. Typical electron-density values are derived, and the study provides an independent confirmation that the electron density and the production rate are related by a square-root law.

**Key words.** Ionosphere (ionospheric disturbances; particle precipitation; polar ionosphere)

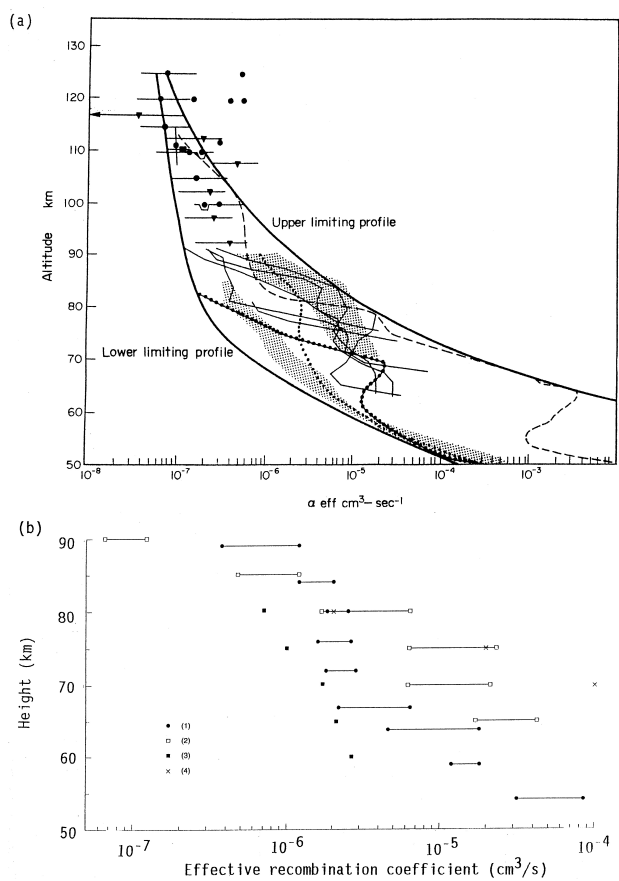
## 1 Background

Polar-cap absorption events (PCA), several good examples of which have occurred during the recent solar maximum, are a direct consequence of energetic protons emitted from an active region of the Sun. On penetrating into the terrestrial atmosphere they enhance the ionization of the mesosphere, which in turn increases the absorption of radio waves in the HF and VHF bands (Bailey, 1959). The incidence and intensity of the event may conveniently be monitored in terms of the radio absorption, using a riometer (Relative Ionospheric Opacity Meter – Little and Leinbach, 1959). The proton fluxes are also routinely monitored above the atmosphere using satellite-borne detectors. An important characteristic of solar proton events is their relative uniformity over the polar regions down to a cut-off latitude at or near 60° geomagnetic latitude (Reid, 1974). The enhancement of electron density may in principle be measured as a function of height by incoherent-scatter radar.

At a given altitude the rate of ion production, which depends on the energy deposited at that level, is determined by the energy spectrum of the proton flux and may be computed from it using an atmospheric model. The ion production rate ( $q$ , in  $\text{cm}^{-3} \cdot \text{s}^{-1}$ ) and the equilibrium electron density ( $N_e$ , in  $\text{cm}^{-3}$ ) at a given height are related by the simple formula

$$q = \alpha_e \cdot N_e^2, \quad (1)$$

where  $\alpha_e$  (in  $\text{cm}^3 \cdot \text{s}^{-1}$ ) is by definition the effective recombination coefficient. The radio absorption at a given altitude

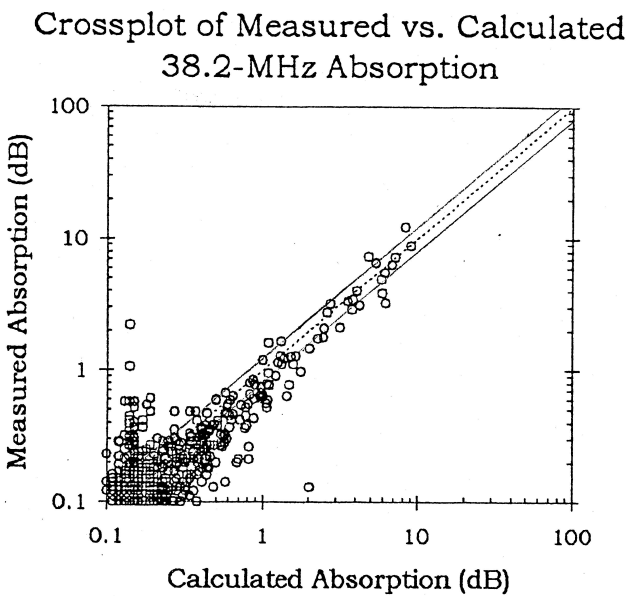


**Fig. 1.** Values of effective recombination coefficient determined by various methods. **(a)** A collection by Penman et al. (1979). **(b)** Recent determinations by incoherent scatter radar (Hunsucker and Hargreaves, 2003). (1) Daytime, summer, the range of values over several days (Reagan and Watt, 1976). (2) Daytime, winter, the range of values over three hours near local noon (Hargreaves et al., 1987). (3) Daytime, spring, afternoon (Hargreaves et al., 1993). (4) Night, spring (Source as 3).

depends on  $N_e \cdot a$ , where “ $a$ ” is the “specific absorption” defined as the absorption per unit path at unity electron density, usually quoted in units of  $\text{dB} \cdot \text{cm}^3 \cdot \text{km}^{-1}$ . Hence, the total absorption over a vertical path is

$$A = \int N_e(h) \cdot a(h) \cdot dh. \tag{2}$$

The specific absorption depends on the effective collision frequency (dominated by the electron-neutral collision frequency  $\nu$ ) and the radio frequency ( $\omega$ ) at which the absorption is measured, in proportion to  $\nu/(\nu^2 + \omega^2)$ . Here,  $\omega$  is the angular frequency ( $=2\pi f$  where  $f$  is in Hz). The specific absorption may be calculated from an atmospheric model. The values used were based on the curves given by Little et al. (1964). No account is taken of possible transitory variations of “ $a$ ” due, for instance, to frictional heating. The riometer technique (in the context of auroral absorption) has been reviewed by Hargreaves (1969).



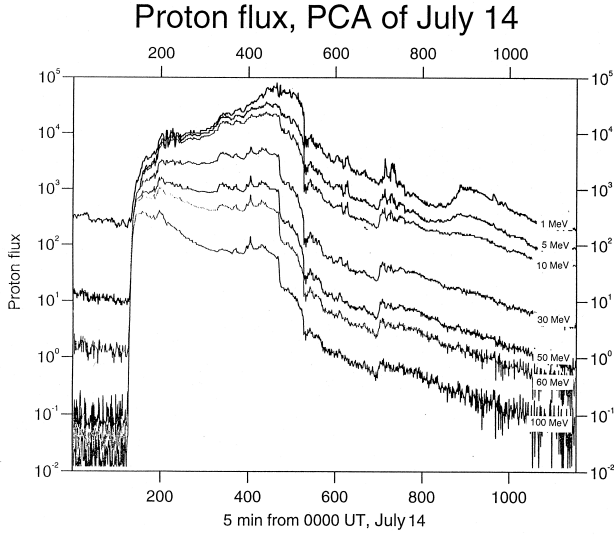
**Fig. 2.** Comparison between measured and calculated absorption during PCAs (Patterson et al., 2001). The outer lines indicate a difference of 20%.

Height (km)	$\sqrt{\text{Production rate}}$	Combined coeff.	Absorption increment
90	$\sqrt{q_{17}}$	$c_{17}$	$\delta A_{17}$
85	$\sqrt{q_{16}}$	$c_{16}$	$\delta A_{16}$
5(n+1)	$\sqrt{q_n}$	$c_n$	$\delta A_n$
15	$\sqrt{q_2}$		$\delta A_2$
10	$\sqrt{q_1}$	$c_1$	$\delta A_1$
$A_{\text{calc}}$			

**Fig. 3.** 5-km slabs used in the computation.

While the radio absorption during a PCA may be measured without difficulty, and values of the proton fluxes are readily available, these do not immediately tell us the effects at a given altitude. Direct measurements of the electron-density profile by rocket or radar are only possible in some instances, and although the incoherent-scatter technique is an attractive one for D-region studies in general, it tends to lack the sensitivity required for PCA studies. The problem may in principle also be solved by using pre-determined values of the effective recombination coefficient. There have been many estimates of this coefficient, but since they vary greatly (Fig. 1) it is not obvious which values should be assumed.

Gledhill (1986) reviewed the literature on  $\alpha_e$  measurements under various circumstances, and derived formulae for its variation with height over the range 50–90 km. The results were given for different types of D-region disturbance, and the PCA results were divided into “day” and “night” according to solar zenith angles less than  $92^\circ$  and greater than  $98^\circ$ , respectively. Recognising the large spread, it was stated



**Fig. 4.** Integral proton fluxes at energies exceeding 1, 5, 10, 30, 50, 60, and 100 MeV during the event which started 14 July 2000, measured by GOES-10 at geosynchronous orbit.

that the results were “not reliable to a factor of 2, and in some cases not even within an order of magnitude”. This qualification succinctly summarises our problem.

Patterson et al. (2001) used Gledhill’s profiles of effective recombination coefficient (day and night) to estimate the total absorption from the observed proton spectrum in a number of events, and obtained reasonable agreement, at least for the daytime. Figure 2 reproduces their comparison between observed and predicted absorption; lines have been added to indicate differences of 20%.

## 2 A new approach

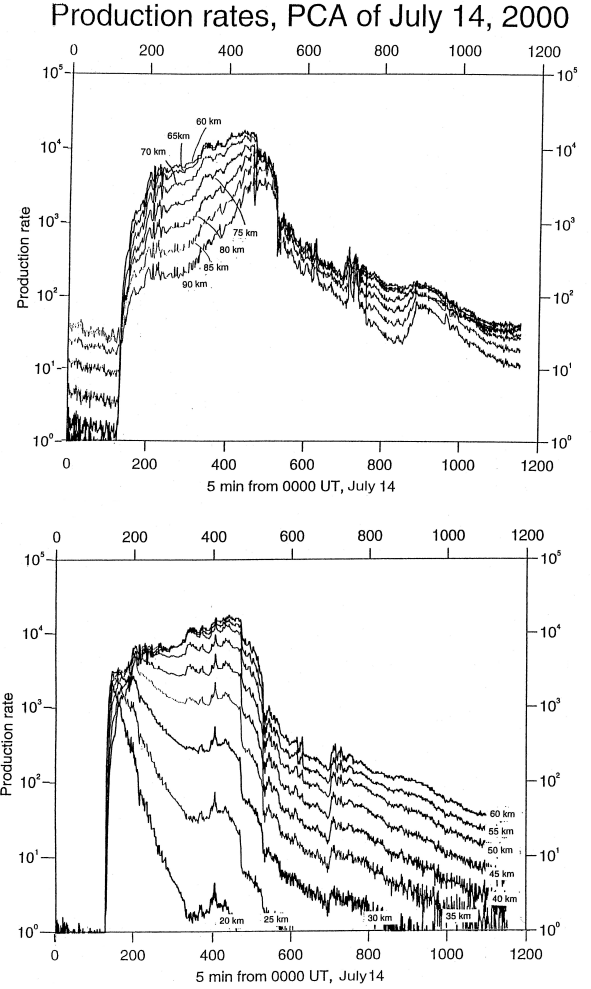
The new approach is as follows. From Eqs. (1) and (2), we can write,

$$A = \int q^{1/2} \cdot c \cdot dh, \quad (3)$$

where

$$c = a/\alpha_e^{1/2}. \quad (4)$$

Given that  $A$  is in decibels (dB),  $h$  in km, and  $a$  and  $\alpha_e$  are in the units stated above, the unit of  $c$  will be  $\text{dB} \cdot \text{cm}^{3/2} \cdot \text{s}^{1/2} \cdot \text{km}^{-1}$ . The ionosphere is divided into 17 slabs of 5 km (Fig. 3). In each slab the incremental absorption is related to the square root of the production rate by the “combined coefficient”,  $c$  (equal to the specific absorption divided by the square root of the effective recombination coefficient). The production rates are computed from geosynchronous-satellite measurements of the proton fluxes by the method of Reid (Reid, 1986; Reid et al., 1991). The scheme of computation is based on laboratory measurements of ionization by



**Fig. 5.** Ion production rates computed for heights from 20 to 90 km during the proton event of July 2000. The curves are displayed in two panels for clarity.

protons in air. Power-law approximations are fitted in sections to the observed fluxes exceeding 1, 5, 10, 30, 50 and 100 MeV, and the energy deposition of individual protons is followed until the particle stops. At selected heights, in this case every 5 km, the deposition is integrated over energy and incidence angle.

The values of “ $c$ ” are varied until the difference between the calculated and measured absorption is a minimum, the criterion being the sum of the squares of the difference taken over many values during the PCA event. We wish to vary  $c(n)$  to minimise the quantity

$$\sum [(A_{\text{calc}} - A_{\text{obs}})/A_{\text{obs}}]^2, \quad (5)$$

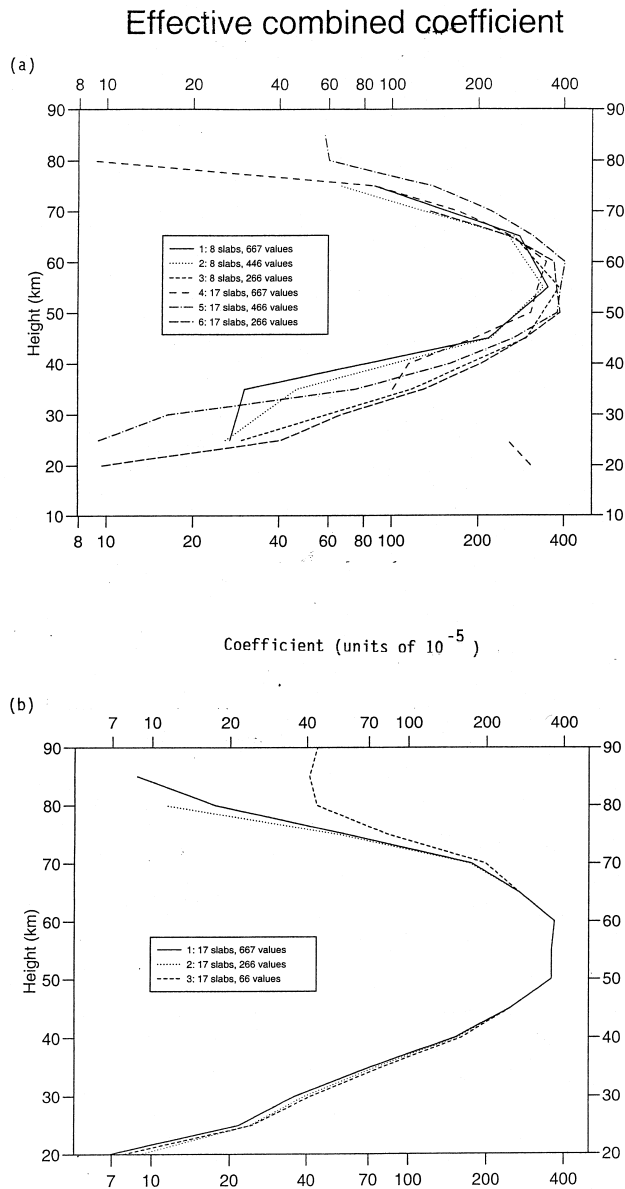
where the summation is over time, and

$$A_{\text{calc}} = 5 \sum q_n^{1/2} \cdot c_n, \quad (6)$$

where this summation is over all the slabs,  $n=1$  to 17.

The mathematical procedure, which is from the NAG Fortran Library, is designed for the solution of an

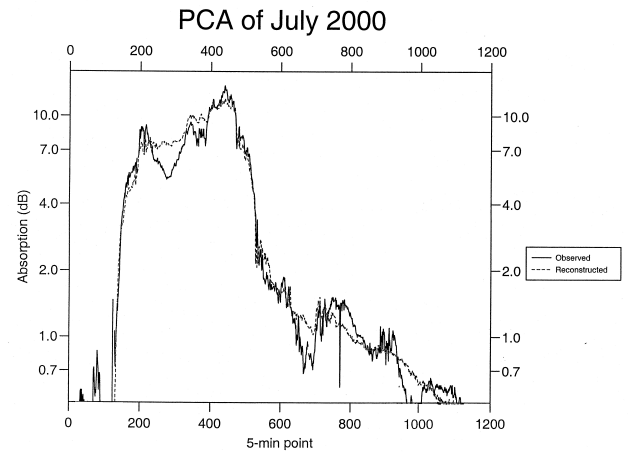




**Fig. 6.** Values of the combined coefficient,  $c$  in units of  $\text{dB} \cdot \text{cm}^{3/2} \cdot \text{s}^{1/2} \cdot \text{km}^{-1}$ , determined for daytime during the event of July 2000. (a) Results from 6 separate runs. (b) Results with some values fixed in advance.

over-determined set of equations (in our case 17 variables but as many as 667 equations). Starting from initial guesses, the program works in (in our case) 17 dimensions and samples with small steps, in order to determine in which direction the minimum lies. A larger step is then taken and the process repeated until a minimum is found. The method relies on the fact that the proton spectrum varies during the event, as a result of which the relative contribution by the various slabs changes with time. No solution would be possible if the spectrum were constant.

Some finesse is required in applying this method. The procedure might find a local minimum which is not the best solution, and there might be mathematical solutions which are



**Fig. 7.** Comparison between observed and reconstructed absorption during the event of July 2000.

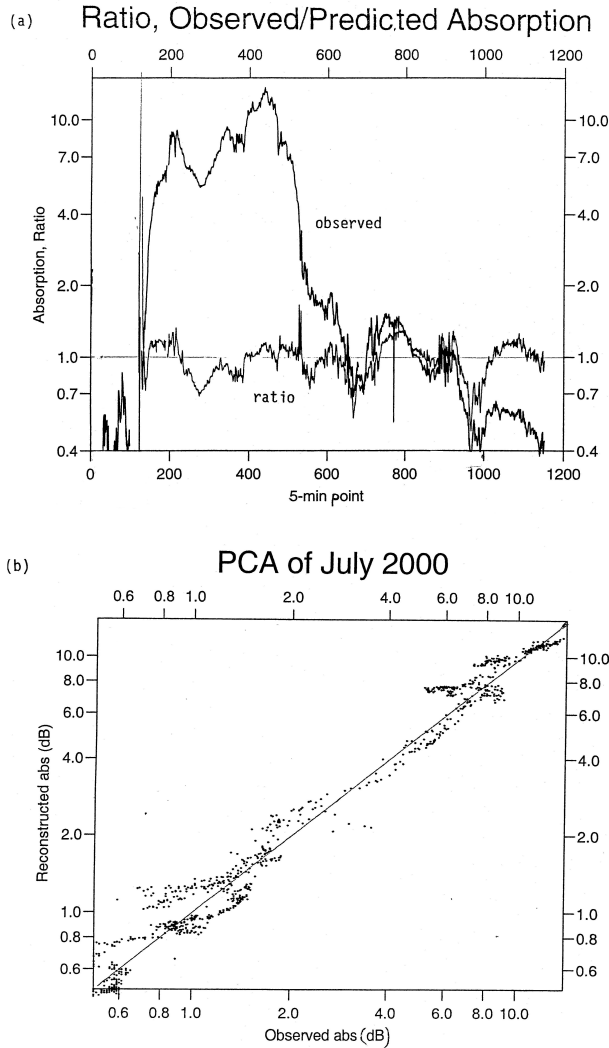
not realistic physically. For instance, we know that " $c$ " cannot be negative and that it should vary smoothly from level to level. Thus, it was found useful to operate on different selections of data and to begin the iteration from different starting points. Moreover, because there will always be noise in the data we cannot expect a perfect solution. We are looking for the best answer that does not contradict existing knowledge and that tends to emerge with some consistency over a number of tests.

### 3 Daytime results

Daytime data are taken from the PCA event of July 2000. The integral proton fluxes above 7 threshold energies (Fig. 4) were measured on a geosynchronous GOES satellite, and production rates from 10 to 90 km were computed from them (Fig. 5). The radio absorption was measured throughout the event with the 38.2-MHz imaging riometer at Kilpisjärvi (Finland), geographic coordinates  $69.05^\circ \text{ N}$ ,  $20.79^\circ \text{ E}$ ,  $L$ -value 5.9, (Browne et al., 1995; Hargreaves and Detrick, 2002), taking the average absorption over the central five beams. This event exceeded 10 dB at its maximum. To maintain accuracy, only values of at least 0.5 dB were used in determining combined coefficients.

Figure 6a shows the values of  $c$  determined using different data selections, and also the results of initial tests run using 8 slabs (of 10 km, centred 15–85 km height). In Fig. 6b, the central values, so determined, have been fixed in the hope of obtaining better precision at the upper and lower extremes. The medians from Fig. 6b have been taken as the working values of  $c$  for this event, and to represent daytime conditions.

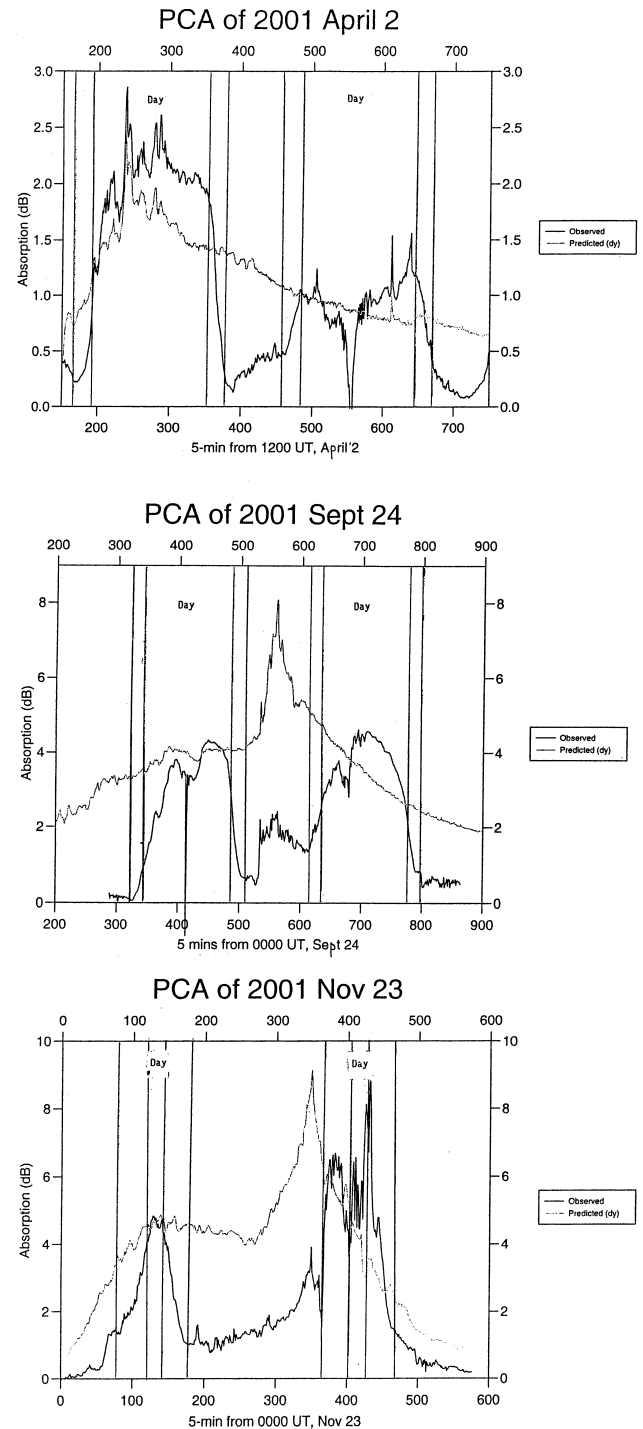
In Fig. 7 the absorption pattern reconstructed using these coefficients is compared with the observed absorption. The ratio of observed/calculated absorption (Fig. 8a) has a median value of 1.0 (horizontal line), and half of the values are within 10% of the median. The measure of agreement



**Fig. 8.** Comparisons between observed and reconstructed absorption, July 2000 all during daytime. **(a)** Ratio of observed to predicted absorption. **(b)** Reconstructed against observed absorption, showing overall linearity.

is substantially the same at all absorption levels (Fig. 8b). This is significant because of the assumption (Eq. 1) that the electron density varies with the square root of the production rate. Were that not the case, Fig. 8b would have deviated from linearity. We therefore have an independent experimental confirmation of the square root law.

In order to test the more general validity of the coefficients, computations and comparisons have been made for the events of 2 April, 24 September and 23 November 2001 (Fig. 9). The vertical lines mark the times when the solar elevation is 0 and  $-10^\circ$ . We take  $0^\circ$  as the limit of “daytime”. The prediction for the April event is low by 30%. The estimate is about right for the daytime sections of the September and November events, particularly on the first day in each case. (Since enhancements of the solar wind tend to reach the Earth later than the increase of proton flux, days after the first



**Fig. 9.** Predicted absorption for the daytime periods of the PCA events of 2001 starting 2 April, 24 September, and 23 November. The vertical lines mark the limits assumed for day and night: solar elevation angles 0 and  $-10^\circ$  respectively. The severe dip on the second day of the April event was caused by a burst of solar radio noise.

are more susceptible to contamination by auroral absorption, whose presence may be suggested by increased irregularity of the riometer trace.)

**Table 1.** Calculated, experimentally determined, and inferred parameters at 40–70 km during daytime PCA events.

Height (km)	40	45	50	55	60	65	70
Combined coefficient (c)	$1.56.10^{-3}$	$2.48.10^{-3}$	$3.58.10^{-3}$	$3.58.10^{-3}$	$3.68.10^{-3}$	$2.68.10^{-3}$	$1.82.10^{-3}$
Specific absorption coeff. (a)	$7.2.10^{-5}$	$7.5.10^{-5}$	$6.1.10^{-5}$	$4.0.10^{-5}$	$2.3.10^{-5}$	$1.15.10^{-5}$	$5.7.10^{-6}$
Effective recomb. ( $\alpha_e$ )	$2.1.10^{-3}$	$9.2.10^{-4}$	$2.9.10^{-4}$	$1.25.10^{-4}$	$3.9.10^{-5}$	$1.85.10^{-5}$	$9.8.10^{-6}$
Event of July 2000							
Incremental absorption (dB/km) at 12:35 UT. Total abs= 12.0 dB	$9.75.10^{-2}$	$2.23.10^{-1}$	$4.13.10^{-1}$	$4.47.10^{-1}$	$4.91.10^{-1}$	$3.57.10^{-1}$	$2.21.10^{-1}$
Fraction of total in slab (total 0.94)	0.04	0.09	0.17	0.19	0.21	0.15	0.09
Electron density $N_e(\text{cm}^{-3})$	$1.4.10^3$	$3.0.10^3$	$6.8.10^3$	$1.1.10^4$	$2.1.10^4$	$3.1.10^4$	$3.9.10^4$
Event of September 2001							
Incremental absorption (dB/km) at 15:35 UT. Total abs= 4.1 dB	$3.85.10^{-2}$	$8.19.10^{-2}$	$1.44.10^{-1}$	$1.53.10^{-1}$	$1.65.10^{-1}$	$1.17.10^{-1}$	$6.91.10^{-2}$
Fraction of total in slab (total 0.95)	0.05	0.10	0.18	0.19	0.20	0.14	0.09
Electron density $N_e(\text{cm}^{-3})$	$5.3.10^2$	$1.1.10^3$	$2.4.10^3$	$3.8.10^3$	$7.2.10^3$	$1.0.10^4$	$1.2.10^4$

Units are c:  $\text{dB.cm}^{3/2}.\text{s}^{1/2}.\text{km}^{-1}$ , a:  $\text{dB.cm}^3.\text{km}^{-1}$ ,  $\alpha_e$ :  $\text{cm}^3.\text{s}^{-1}$

#### 4 Some inferences: incremental absorption, recombination coefficient and electron density

In the July event, the incremental absorption (defined as the absorption in dB/km) comes to a maximum at 60 km during most of the event (Fig. 10), though the contributions at 55 and 50 km are as large during the early part, and the absorbing layer extends well below even those levels during the growth phase (when the proton spectrum is hardest). Generally, above 70 km and below 40 km the contribution to the absorption (per km) is less than 10% of that at 60 km. The slabs centred from 45 to 65 km account for 80% of the total absorption. Less than 1% of the total absorption arises below 30 km or above 80 km.

From the definition of the “combined coefficient”, c, and since we know the “specific absorption”, a, as a function of altitude, we may compute values of the “effective recombination coefficient”. The most accurate values will be those for the slabs that make the largest contribution to the total absorption, i.e. from 45 to 65 km. Figure 11 compares those values (extended by one slab at each end) with Gledhill’s (1986) daytime profile of this quantity. Our values are larger by a factor of 1.5 to 2. Recall that the Gledhill values were obtained by reviewing the existing literature, in which the values show a large spread. It is interesting that the plots of Fig. 11 show similar gradients over their common height range. The new determination is entirely independent of previous estimates and, moreover, it extends down to altitudes for which other data are sparse.

Specimen values of the electron density during PCA events may also be estimated from Eqs. 1, 2 and 4. Table 1 gives such values for the height range 40–70 km, for times near the daytime maxima of the events of July 2000 and September 2001. These values are consistent with direct measurements (by incoherent-scatter radar – Fig. 1b) at

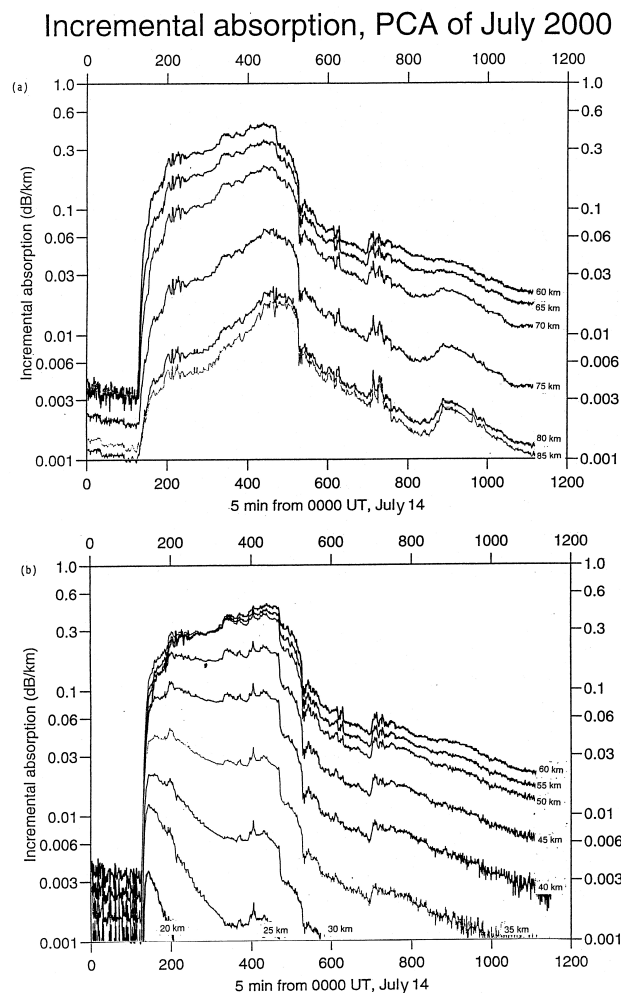
the higher altitudes. However, no comparable radar data are believed to exist for the lower levels.

Table 1 also gives the values of combined coefficient, specific absorption coefficient, and effective recombination coefficient, all for daytime conditions as derived from the event of July 2000.

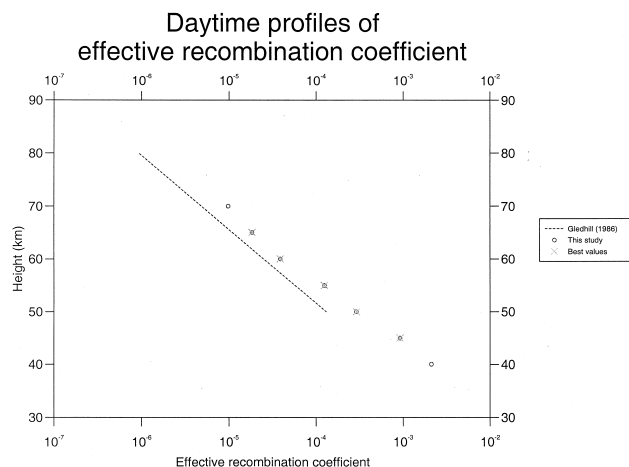
#### 5 Night-time results

We expect the results to be rather different for the night periods, when free electrons are removed at the lower altitudes by attachment to neutral species. Night periods were selected from the events of 23 November 2001 and 24 September 2001, the criterion for “night” being that the Sun was more than  $10^\circ$  below the horizon, this being the depression angle for which riometers generally indicate that the transition to night-time absorption levels is complete. Data from 2 April were not included because the absorption values fell below 0.5 dB. To obtain the greatest range of the proton spectrum, we wish to include observations early and late in the event. The first phase, however, tends to be of short duration, whereas the last will be more subject to auroral contamination. This is a greater problem by night than by day because the component of absorption due to the protons is relatively smaller by night for a given proton flux.

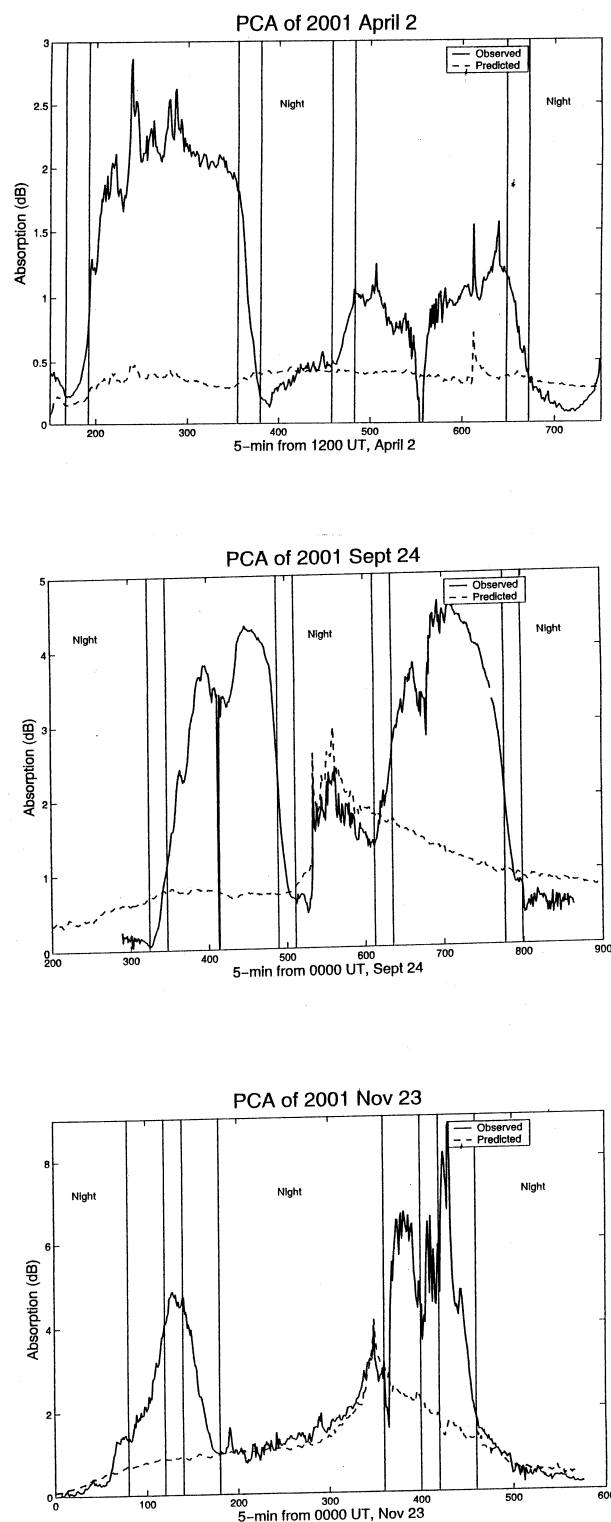
It was indeed found that the consistency of results from different runs of the procedure was poorer by night than had been the case by day. Taking what is considered to be “a good estimate”, Fig. 12 shows the reconstruction of night periods for the events of April, September and November, compared with the observations. The November event is reproduced quite well, as is the September event except for the dip immediately after sunset. The sharp enhancement



**Fig. 10.** The incremental absorption (dB/km) at heights 20 to 85 km, during the PCA of July 2000, showing the contribution maximising at 50–65 km during most of the event.



**Fig. 11.** Effective recombination coefficients for daytime derived by the present method, compared with Gledhill's (1986) values.



**Fig. 12.** Predicted absorption for the night periods of the 2002 events beginning 2 April, 24 September, and 23 November. The vertical lines are as on Fig. 9.

**Table 2.** Values of  $c$  derived from various night-time runs. **(a)** From three different selections of data: (i) One period in November; (ii) Periods from November and September; (iii) Three periods from November with a shorter one from September. Three progressive stages during run (iii) are shown. **(b)** Progressive stages of a refining run. The data set is the same as (iii), but the starting values are based on an average of the results in Table (a). **(c)** Further tests in which only a limited number of levels were permitted. In (v), only 75 and 80 km were allowed to be non-zero. Four levels were allowed in (vi). In each case the data set is the same as in (iii) and (iv), and the starting guess is from the bottom line of Table 2b.

(a) Values of c from selection of night-time data (first attempts)								
Data	No.	r.m.s. residual	Height (km)					
	of points		65	70	75	80	85	90
(i)	181	0.114	0	0	$1.44.10^{-3}$	$3.22.10^{-3}$	0	0
(ii)	281	0.185	0	$5.75.10^{-4}$	$4.08.10^{-3}$	0	0	0
(iii)	257	0.235	0	0	$4.33.10^{-3}$	0	0	$1.56.10^{-3}$
		0.235	0	0	$4.07.10^{-3}$	$1.64.10^{-5}$	$1.56.10^{-3}$	$1.167.10^{-5}$
		0.235	0	0	$4.16.10^{-3}$	0	$1.41.10^{-3}$	0
(b) Values of c from night-time data (refining run)								
Data	No.	r.m.s. residual	Height (km)					
	of points		65	70	75	80	85	90
(iv)	257	0.238	0	$1.15.10^{-4}$	$2.87.10^{-3}$	$1.69.10^{-3}$	$5.94.10^{-4}$	$3.15.10^{-4}$
		0.236	0	$1.47.10^{-4}$	$2.91.10^{-3}$	$1.74.10^{-3}$	$6.50.10^{-4}$	$3.87.10^{-4}$
		0.236	$2.81.10^{-6}$	$1.47.10^{-4}$	$2.91.10^{-3}$	$1.74.10^{-3}$	$6.50.10^{-4}$	$3.87.10^{-4}$
		0.235	$2.30.10^{-5}$	$2.10.10^{-4}$	$2.97.10^{-3}$	$1.72.10^{-3}$	$5.51.10^{-4}$	$1.83.10^{-4}$
		0.235	0	$2.21.10^{-4}$	$2.99.10^{-3}$	$1.72.10^{-3}$	$5.43.10^{-4}$	$1.65.10^{-4}$
		0.235	0	0	$3.42.10^{-3}$	$1.66.10^{-3}$	$3.90.10^{-4}$	$1.27.10^{-4}$
		0.235	0	0	$3.46.10^{-3}$	$1.66.10^{-3}$	$3.66.10^{-4}$	$9.42.10^{-5}$
(c) Values of c from night-time data (further tests)								
Data and conditions	No. of points	r.m.s. residual	Height (km)					
			65	70	75	80	85	90
(v)	257	0.235	–	–	$3.29.10^{-3}$	$2.20.10^{-3}$	–	–
(vi)	257	0.235	–	$2.28.10^{-4}$	$2.97.10^{-3}$	$1.77.10^{-3}$	$6.53.10^{-4}$	–
		0.235	–	0	$3.27.10^{-3}$	$1.75.10^{-3}$	$6.22.10^{-4}$	–

during the September event was due to an enhancement of proton fluxes at the lower energies due to the arrival of a solar-wind shock, whose effect appeared only in the production rates above 75 km. The shape of the variation by night is poorly reproduced in the April event, though the absolute difference is no more than 0.3 dB.

The values of  $c$  obtained from various runs for the night periods are shown in Table 2a. The residual r.m.s. error varies only slightly between these cases, being close to 23.5%, yet the “ $c$ ” values which the procedure deems to give the best fit vary widely. Evidently, the result is very sensitive to small variations in the data. On the other hand, this also means that there are several sets of “ $c$ ” which will predict the absorption from the proton flux equally well, presumably because

variations in the production rate are well correlated over the effective height range. Note that  $c=0$  below 70 km in every run, which accords with the well established view that the effective recombination coefficient increases sharply at sunset at the lower altitudes. The present result says that at night the radio absorption cuts off sharply below the 70 km slab and that any absorption contributions lower down are negligible in relation to the total. There is experimental evidence for a sharp increase in the ratio of negative ions to electrons below 75 km (Hall et al., 1988).

The first row of Table 2b is an average of the values in Table 2a, adjusted to smooth the transition between 75 and 85 km. This was then used as the starting guess and the procedure was run again, giving successively the results in

**Table 3.** Distribution of absorption with height at selected times in the November and September events, using coefficients from Tables 2b and c.

Set of c	h (km)	Incremental absorption (dB/km)								% of total absorption in slab
		23 November event				24 September event				
		Point				Point				
Table 2b, bottom row	UT	60	200	340	520	300	520	560	830	
		05:00	16:40	04:20	19:20	01:00	18:20	22:40	21:10	
	75	.075	.144	.335	.064	.081	.131	.381	.109	63–71
	80	.027	.050	.132	.029	.030	.048	.169	.051	25–39
	85	.004	.008	.023	.006	.005	.008	.033	.010	4–6
	90	.001	.001	.005	.000	.001	.001	.007	.002	0–1
Calculated total		.54	1.02	2.48	.50	.58	.94	2.95	.86	
Observed total		.52	.98	2.70	.34	.19	.71	2.34	.54	
h (km)		Incremental absorption								
Table 2c bottom row	75	.072	.135	.317	.060	.077	.124	.360	.103	59–67
	80	.029	.053	.139	.031	.031	.050	.179	.055	26–31
	85	.007	.013	.039	.010	.008	.013	.056	.017	6–10
Calculated total		.54	1.01	2.48	.51	.58	.94	2.98	.88	
Observed total		.52	.98	2.70	.34	.19	.71	2.34	.54	

**Table 4.** Estimated electron densities ( $\text{cm}^{-3}$ ) at 75 and 80 km for the events and times used in Table 3.

Set of c	h (km)	Specific absorption (a)	Electron density ( $\text{cm}^{-3}$ )							
			23 November event				24 September event			
			Point				Point			
			60	200	340	520	300	520	560	830
Table 2b, bottom row	75	$2.65 \cdot 10^{-6}$	$2.9 \cdot 10^4$	$5.4 \cdot 10^4$	$1.3 \cdot 10^5$	$2.4 \cdot 10^4$	$3.1 \cdot 10^4$	$4.9 \cdot 10^4$	$1.4 \cdot 10^5$	$4.1 \cdot 10^4$
	80	$1.2 \cdot 10^{-6}$	$2.3 \cdot 10^4$	$4.2 \cdot 10^4$	$1.1 \cdot 10^5$	$2.4 \cdot 10^4$	$2.5 \cdot 10^4$	$4.0 \cdot 10^4$	$1.4 \cdot 10^5$	$4.3 \cdot 10^4$
Table 2c, bottom row	75	$2.65 \cdot 10^{-6}$	$2.7 \cdot 10^4$	$5.1 \cdot 10^4$	$1.2 \cdot 10^5$	$2.3 \cdot 10^4$	$2.9 \cdot 10^4$	$4.7 \cdot 10^4$	$1.4 \cdot 10^5$	$3.9 \cdot 10^4$
	80	$1.2 \cdot 10^{-6}$	$2.4 \cdot 10^4$	$4.4 \cdot 10^4$	$1.2 \cdot 10^5$	$2.6 \cdot 10^4$	$2.6 \cdot 10^4$	$4.2 \cdot 10^4$	$1.5 \cdot 10^5$	$4.6 \cdot 10^4$

Table 2b, the final row of which was taken as the working values for “night” in deriving the predictions in Fig. 12. Note that the change in the residual error is very slight during this run.

As a further test (Table 2c), the values of  $c$  obtained above were taken as the starting guess and only 2 levels (75 and 80 km) were allowed to generate non-zero values of  $c$ . This was repeated, allowing contributions from only 4 levels (70 to 85 km). Table 2c gives the result at 2 stages of this second run. Note that the value of  $c$  at 70 km generated in this run is either very small or zero. It was shown above that at night there are no significant contributions to the absorption from below 70 km. We can now add that when contributions are restricted to the slabs from 70 to 90 km, there are virtually no contributions from 70 or 90 km either. The night-time absorption is effectively confined to the three slabs centred from 75 to 85 km. Note that the r.m.s. residual is virtually constant in all the results shown here. Any of these sets of

coefficients will give virtually the same accuracy of PCA prediction (i.e. 23–24%).

## 6 Inferences for the night period

Table 3 gives the incremental absorption and the distribution of absorption with height for selected points (all at night) during the November and September events. The points labelled November: points 60, 200, 340 and 520; and September: points 300, 520, 560 and 830, are, respectively, at 23 November, 05:00 UT and 16:40 UT, 24 November, 04:20 UT and 19:20 UT, 25 September 01:00 UT, 18:20 UT and 22:40 UT, and 26 September, 21:10 UT. The  $c$  values used are the final values from Table 2b, and from Table 2c, respectively. Clearly, the major part of the night-time absorption during these PCAs comes from the 75 and 80 km slabs, the former producing 60–70% of the total absorption, and the latter 25–30%. The contribution of the 85-km slab is a minor one at

4–10% (typically in the late and early phases of the event respectively).

When it comes to estimating the effective recombination coefficients, we should, as for the daytime, only use values from slabs that make a major contribution to the total absorption. For night, this restricts us to 75 and 80 km. Most of the values of  $c$  for 75 and 80 km are in the range  $2.91 \cdot 10^{-3}$  to  $4.33 \cdot 10^{-3}$  and  $1.66 \cdot 10^{-3}$  to  $2.20 \cdot 10^{-3}$ , respectively. Following the method of Sect. 4, these lead to values of effective recombination coefficient of  $8.3 \cdot 10^{-7}$  to  $3.7 \cdot 10^{-7}$  and  $5.2 \cdot 10^{-7}$  to  $3.0 \cdot 10^{-7}$ . These values are smaller than those from the Gledhill (1986) formula for night ( $1.56 \cdot 10^{-5}$  and  $4.83 \cdot 10^{-6}$ ), and are close to the “lower limiting values” in Fig. 1 though still within the range shown there.

The values of electron density at 75 and 80 km corresponding to the 8 samples detailed in Table 3 are given in Table 4. We note that at these heights the electron density values are similar to each other, being about 4 to 5 times  $10^4$  if the total absorption is 1 dB, with approximate proportionality between electron density and absorption.

## 7 Conclusions and discussion

The method outlined in the paper uses the variations of calculated ion production rate as a function of height in comparison with the total radio absorption measured with a riometer, to estimate the magnitude of the absorption contribution as a function of height. It has been applied to recent polar-cap absorption events, and to day and night periods separately.

Working with 5-km slabs it was found that most of the absorption occurs in the slabs centred from 45 to 65 km by day (80%), and at 75 to 80 km by night. By day, typically, less than 1% of the absorption arises above 80 km or below 30 km. At night there is no significant contribution to the absorption below the 75-km slab or at 90 km.

The observed linearity between the observed and computed values of total absorption provides an experimental confirmation of the square-law relationship between production rate and electron density (Eq. 1).

The method does not require values of the effective recombination coefficient to be assumed, but estimates of this quantity may be obtained from the results. The daytime values so deduced are comparable with the values recommended in Gledhill's (1986) literature review, though they also extend the range down to 40 km. Night values are considerably smaller than Gledhill's at 75 and 80 km. Electron density values may also be deduced for the heights contributing the most absorption, and these are generally comparable with results from the limited incoherent-scatter radar data available. Values derived in this manner may be useful estimates at times when direct measurements are not available.

This analysis assumes that the values of the “combined coefficient”,  $c$ , are constant at a given height. In fact, there is some evidence from radar data that the effective recombination coefficient may vary slowly with time (Reagan and Watt, 1976; Hargreaves et al., 1993), and such a variation

would plainly reduce the accuracy of the prediction. Further, no seasonal variations have been taken into account in the present study. The method takes no account of any variations in the specific absorption coefficient (Eq. 2), nor of auroral electrons which are most likely to arrive after the first day of the proton event and have most effect at the higher altitudes, therefore making their greatest relative contribution at night. These caveats considered, the overall accuracy of 20–25% which the method achieves appears reasonable and worthwhile.

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